A Critical View of ISM Modelling in Cosmological Simulations
Supernova Feedback

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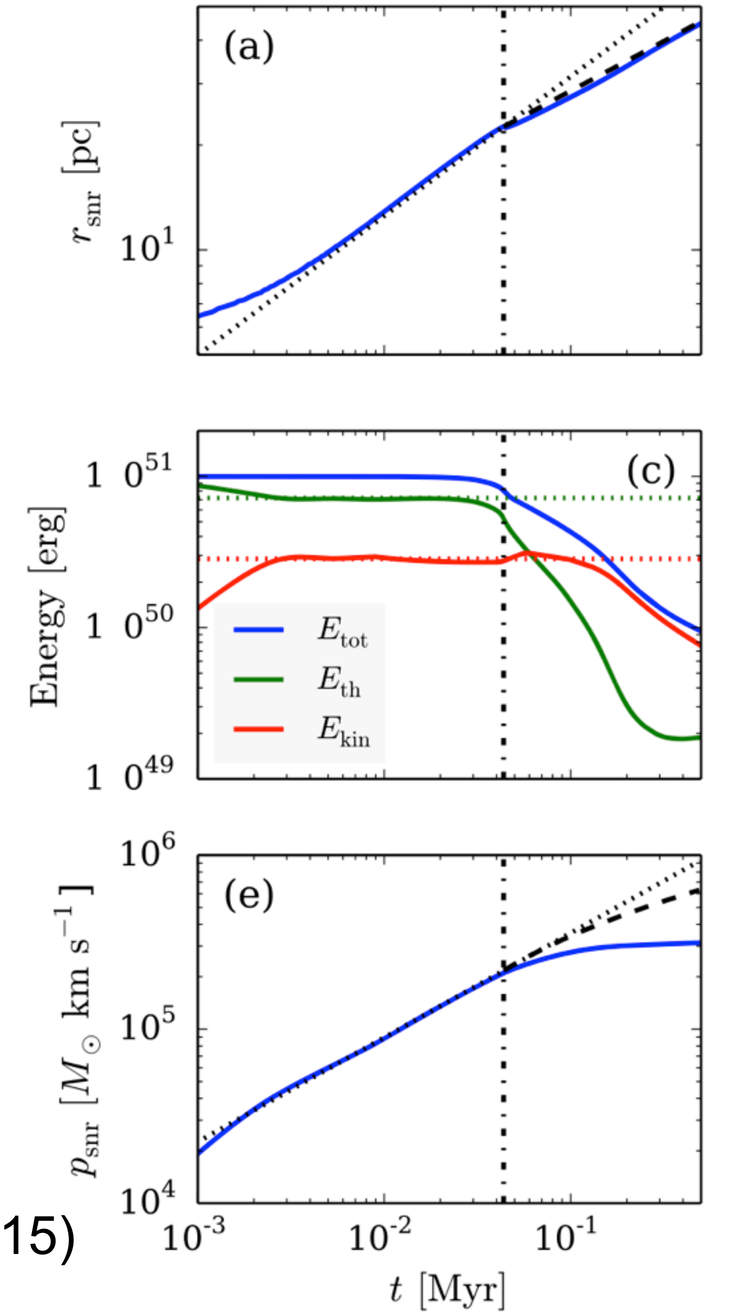
Supernova remnant evolution

- Free expansion (~10^2 yrs)
- **Sedov-Taylor** (~10^4 yrs)
- Pressure driven snowplough (~10^6 yrs)
- Momentum driven snowplough, remnant merges with ambient medium



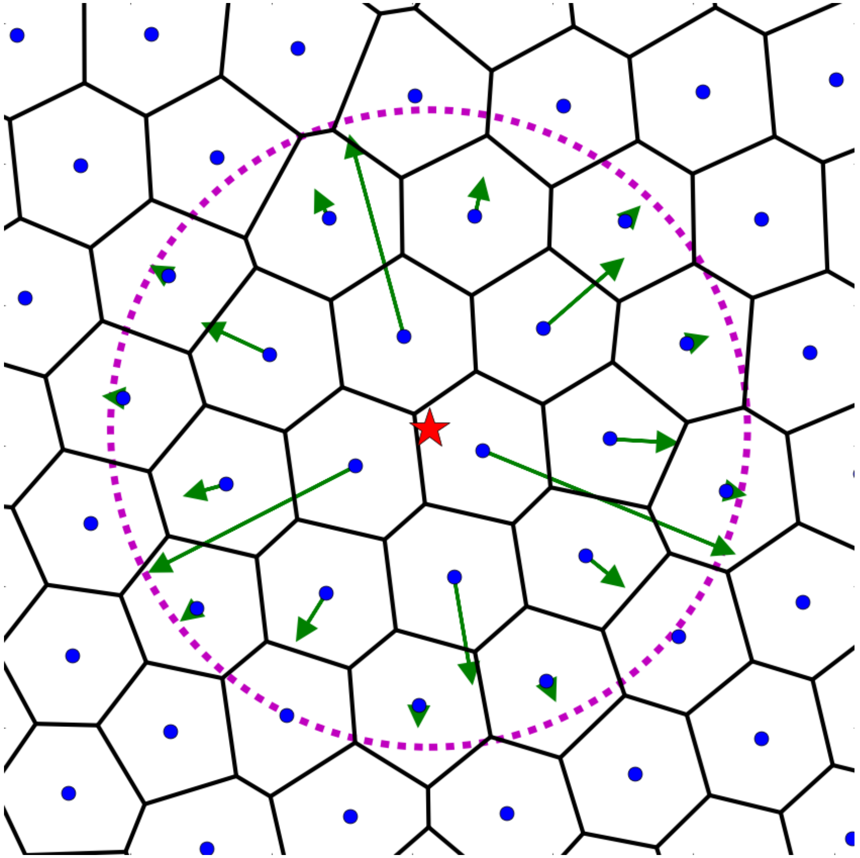
(Kim & Ostriker 2015)

Numerical Scheme

- Moving mesh code AREPO (Springel 2010)

$$\dot{\rho}_* = \frac{\varepsilon \rho}{t_{\rm ff}} \propto \rho^{3/2}, \;\; \rho > \rho_{\rm SF}$$

- SN rates as function of age and metallicity from STARBURST99 (Leitherer+ 1999)
- Stochastically trigger SN events
- Inject mass, energy and/or momentum to surrounding gas, weighted by kernel.

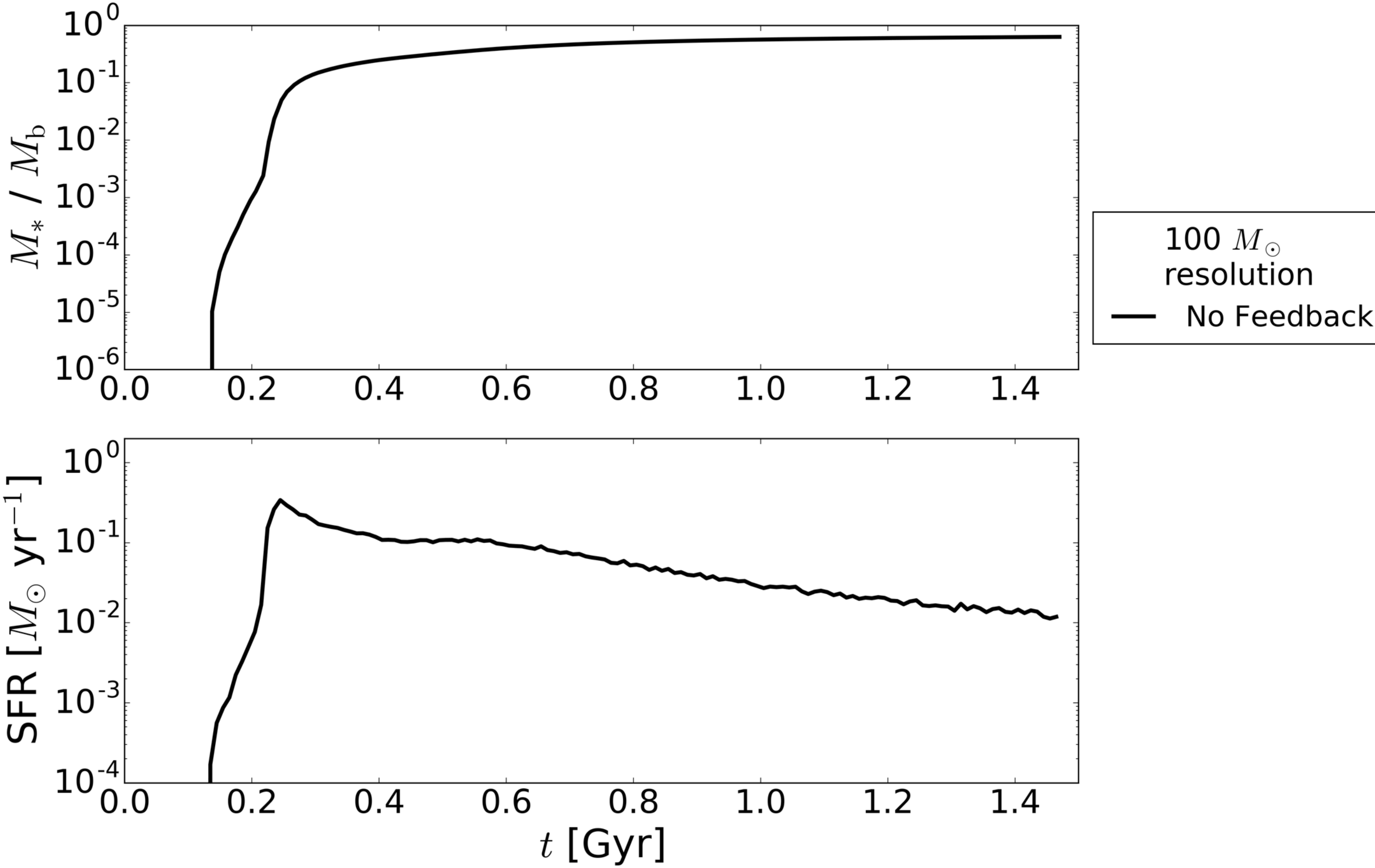


Test Case: Radially Collapsing Gas

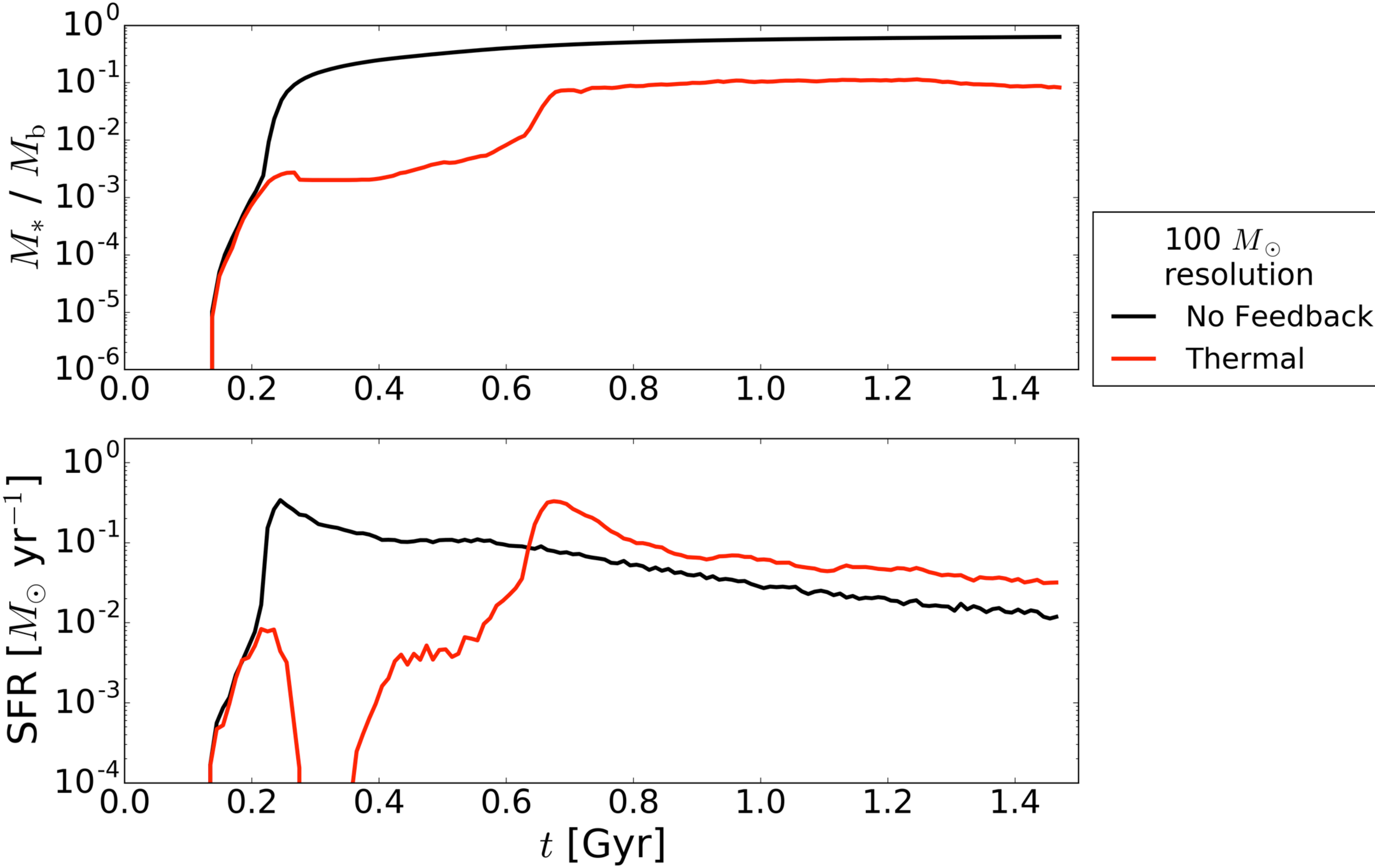
- $2.3\times10^9\ M_\odot$ dark matter (static potential), $1.35\times10^8\ M_\odot$ gas
- Cored isothermal profile with core radius $r_c = 0.5$ kpc

$$\rho(r) = \frac{\rho_0}{1+(r/r_c)^2}$$

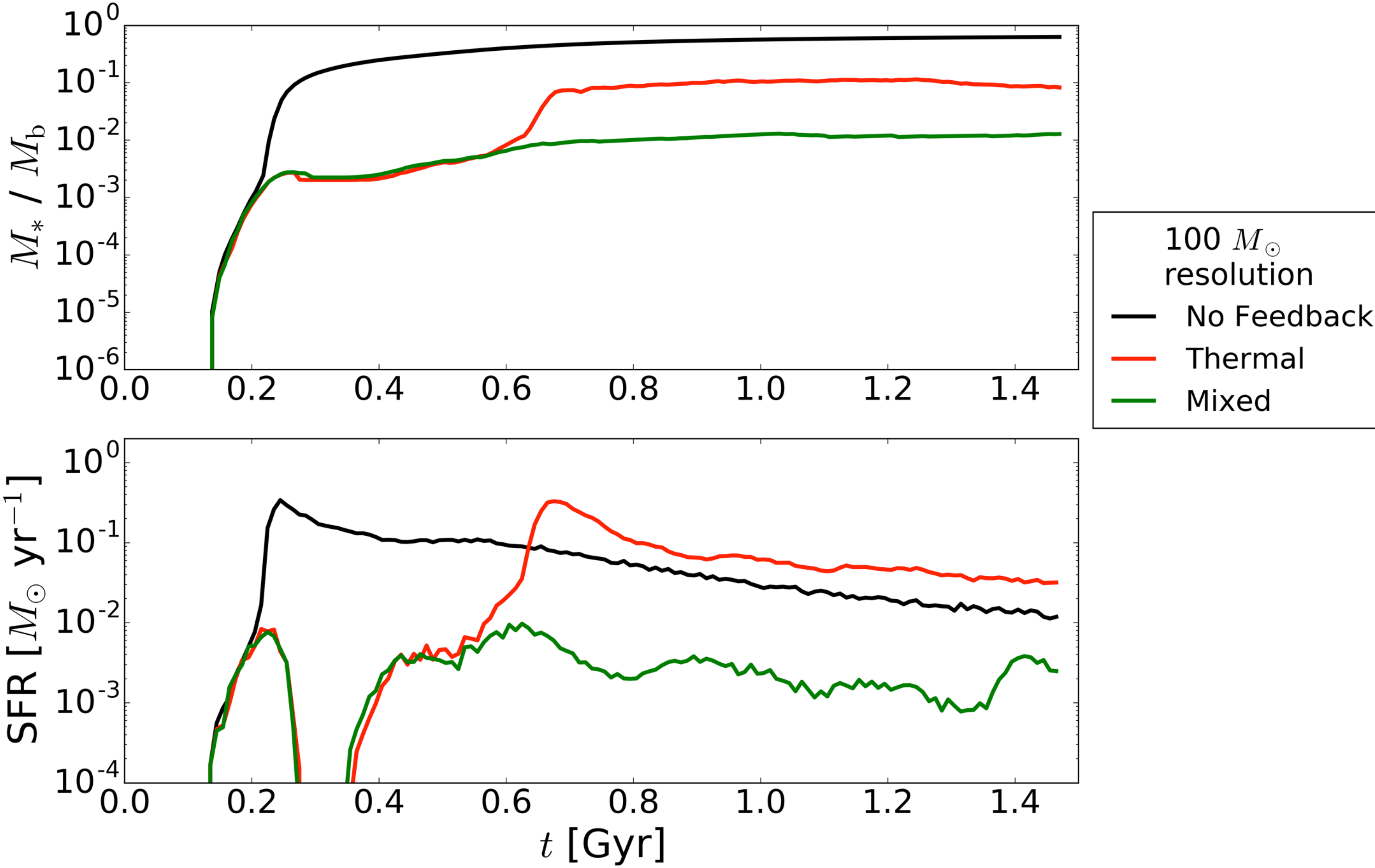
- Gas initially static, $T = 1.12\times10^4$ K, provides 1/3 thermal support required for equilibrium
- Primordial atomic cooling



$M_* \,/\, M_\mathrm{b}$
10^0
10^{-1}
10^{-2}
10^{-3}
10^{-4}
10^{-5}
10^{-6}
0.0
0.2
0.4
0.6
0.8
1.0
1.2
1.4
100 $M_\odot$ resolution
No Feedback
SFR [$M_\odot$ yr^{-1}]
t [Gyr]



$M_* \ / \ M_\mathrm{b}$
SFR [$M_\odot$ yr^{-1}]
t [Gyr]
100 $M_\odot$ resolution
No Feedback
Thermal



10^0
10^{-1}
10^{-2}
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$M_* \ / \ M_\mathrm{b}$
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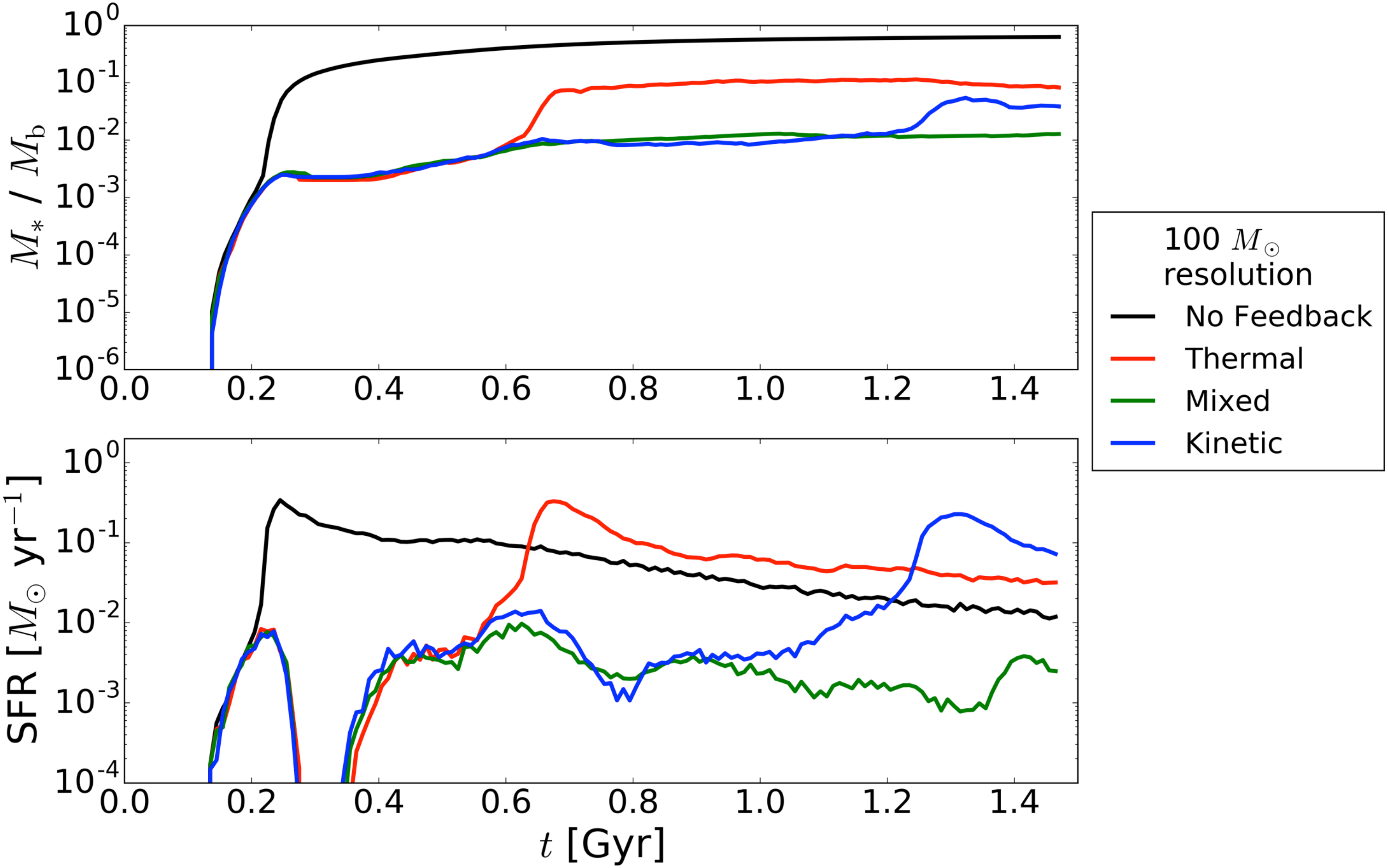
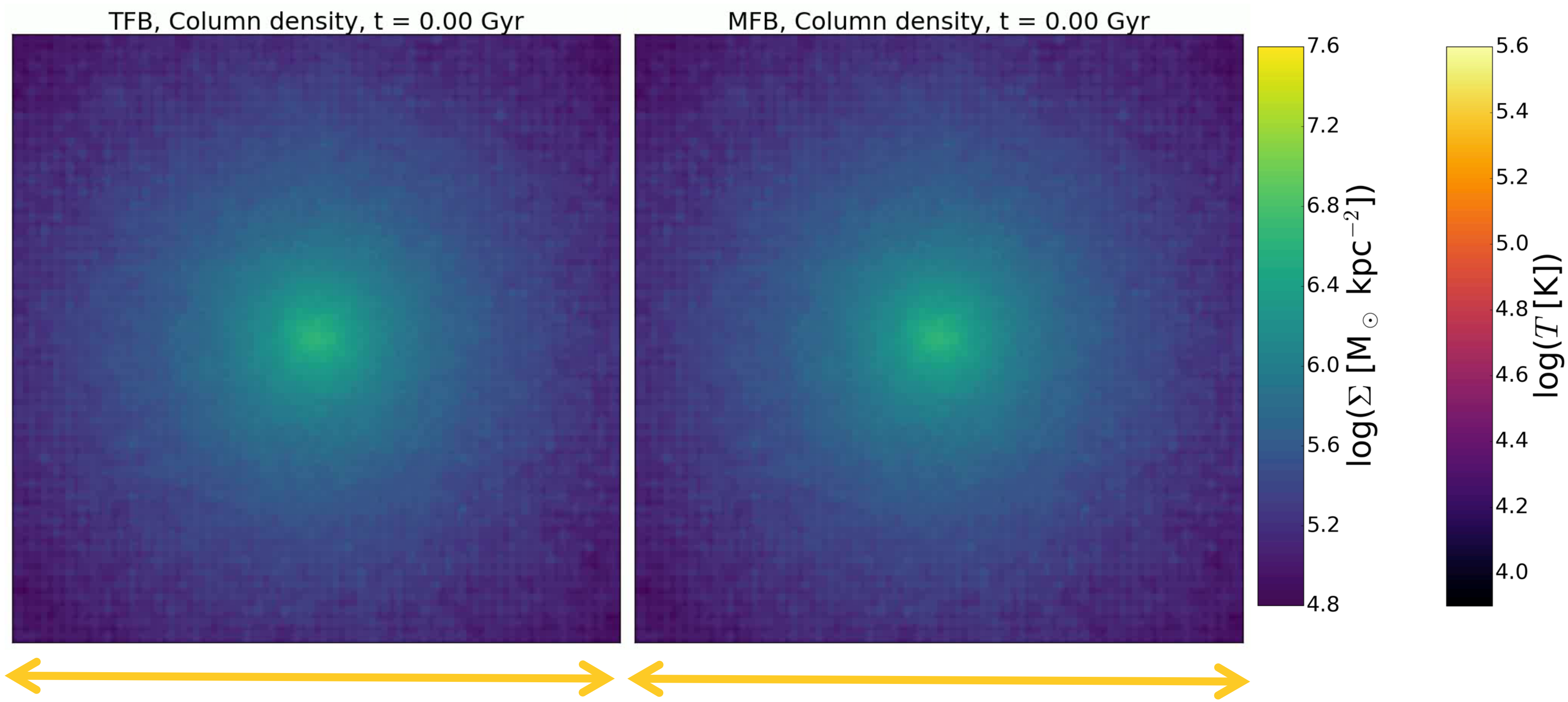
1.0
1.2
1.4
100 $M_\odot$ resolution
No Feedback
Thermal
Mixed
SFR [$M_\odot$ yr^{-1}]
t [Gyr]



M_* / M_b
SFR [$M_\odot$ yr^{-1}]
t [Gyr]
100 $M_\odot$ resolution
No Feedback
Thermal
Mixed
Kinetic

Thermal   Mixed



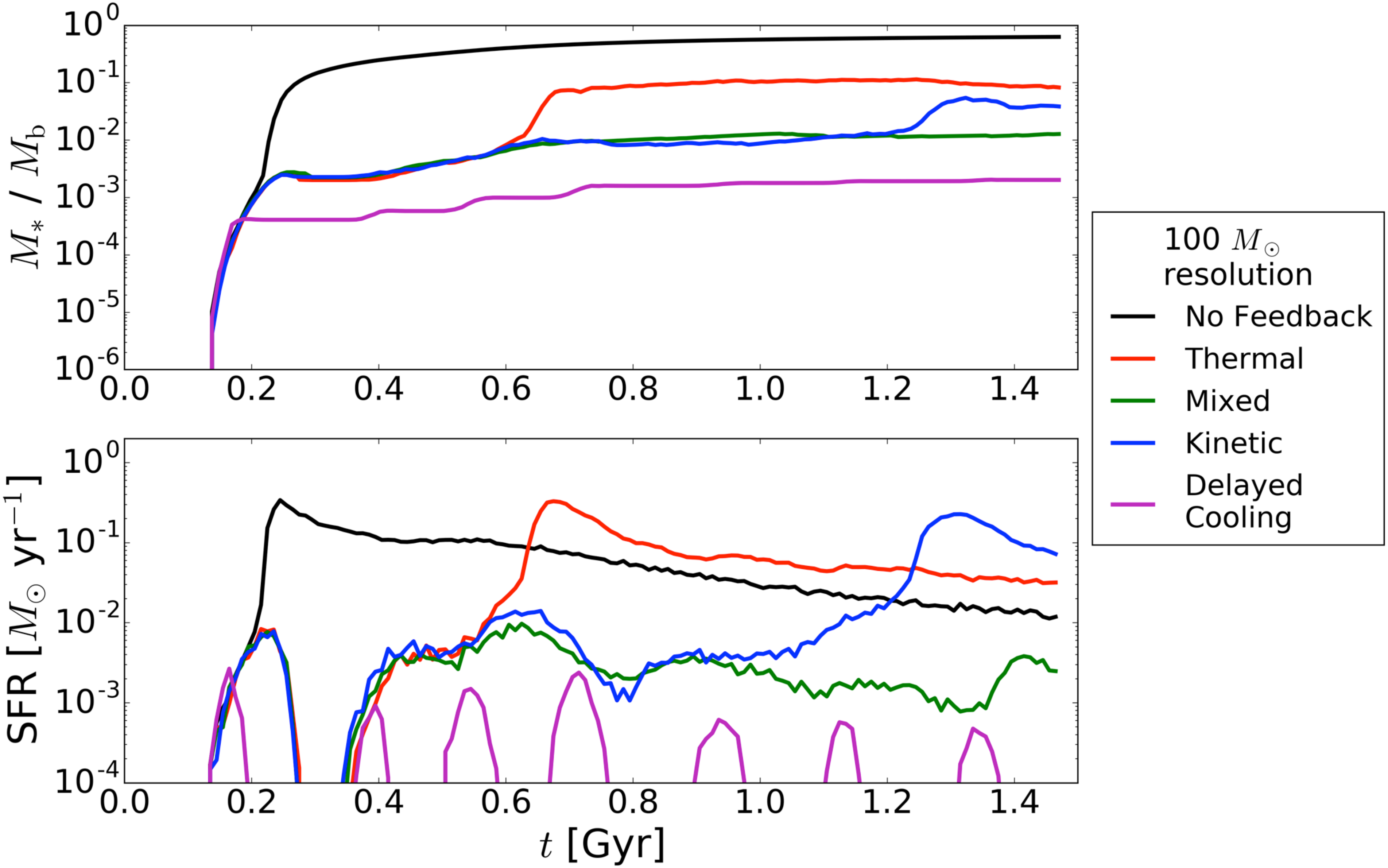
15 kpc   15 kpc

Delayed Cooling

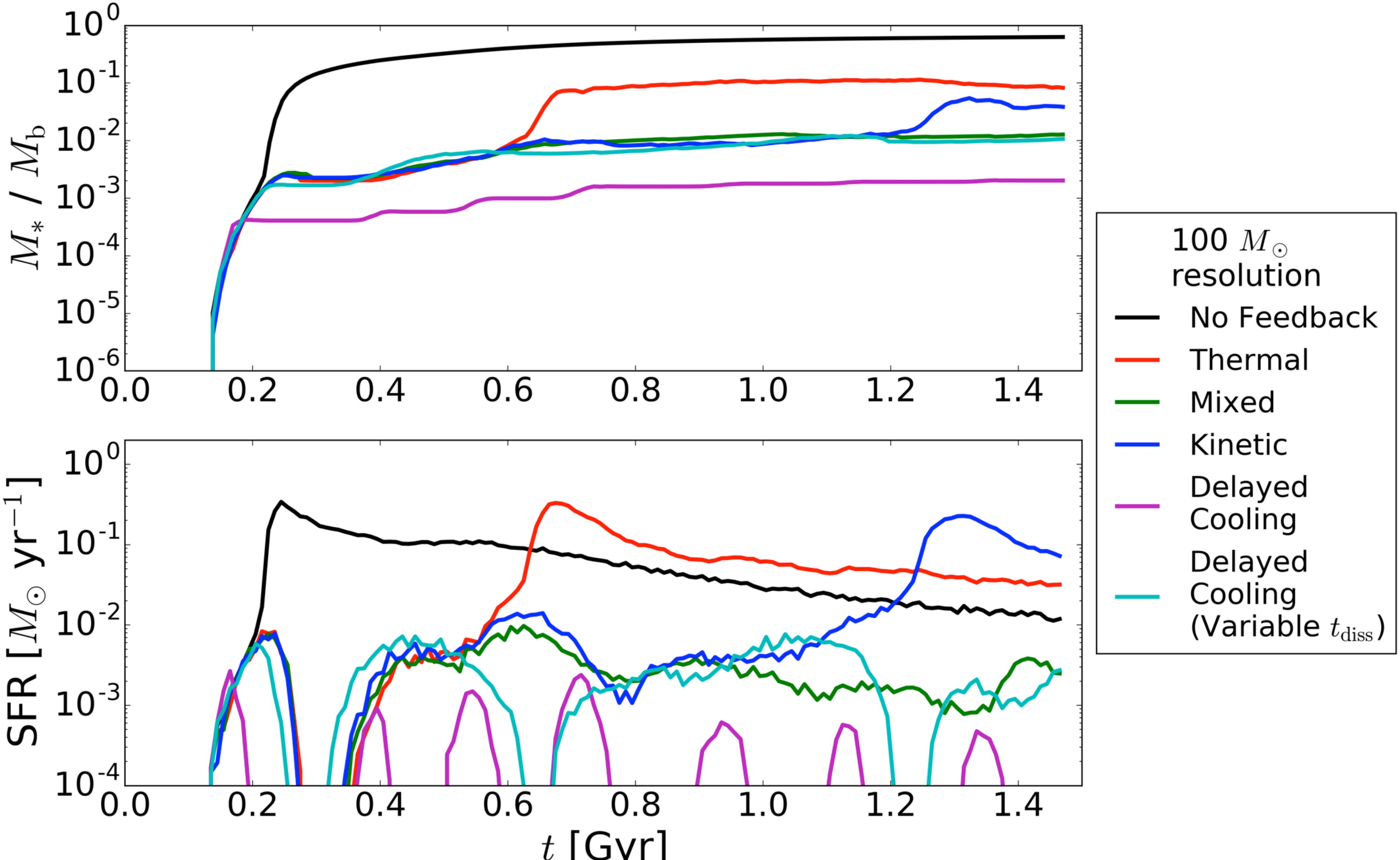
- Delayed Cooling Feedback (based on Teyssier+ 2013)
- Assume E_{SN} locked up in non-thermal processes with much longer dissipation time than cooling time.
- Inject thermal energy but track injected SN energy in passive scalar ε_{SN}
- Advect with gas and dissipate as $\dot{\varepsilon}_{SN} = \varepsilon_{SN} / t_{\mathrm{diss}}$
- Turn off radiative cooling when:

$$\sigma_{\mathrm{SN}} = \sqrt{2\varepsilon_{\mathrm{SN}}} > 10 \ \mathrm{km\ s^{-1}}$$

- $t_{\mathrm{diss}} = 10 \ \mathrm{Myr}$ or $t_{\mathrm{diss}} = \dfrac{\Delta x}{\sigma_{\mathrm{SN}}}$



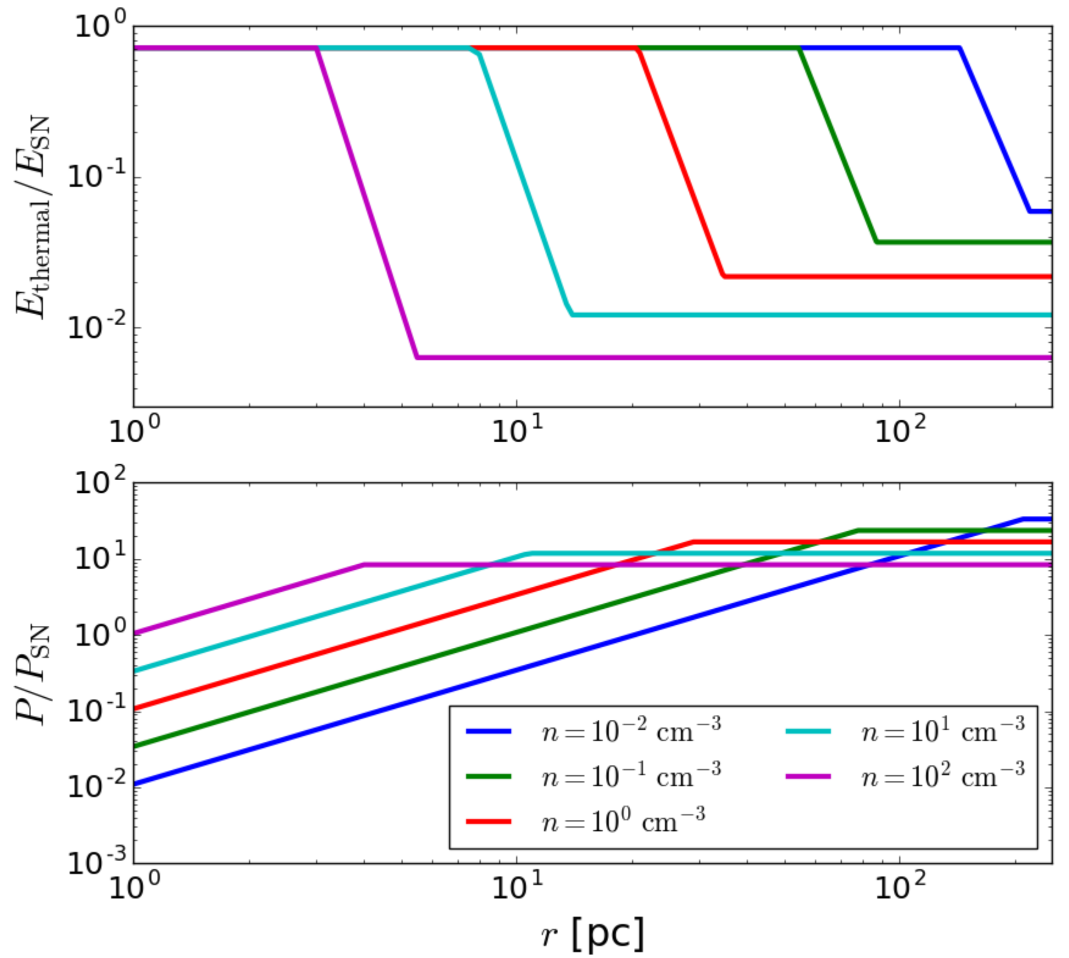
M_* / M_b
SFR [$M_\odot$ yr^{-1}]
t [Gyr]
100 $M_\odot$ resolution
No Feedback
Thermal
Mixed
Kinetic
Delayed Cooling

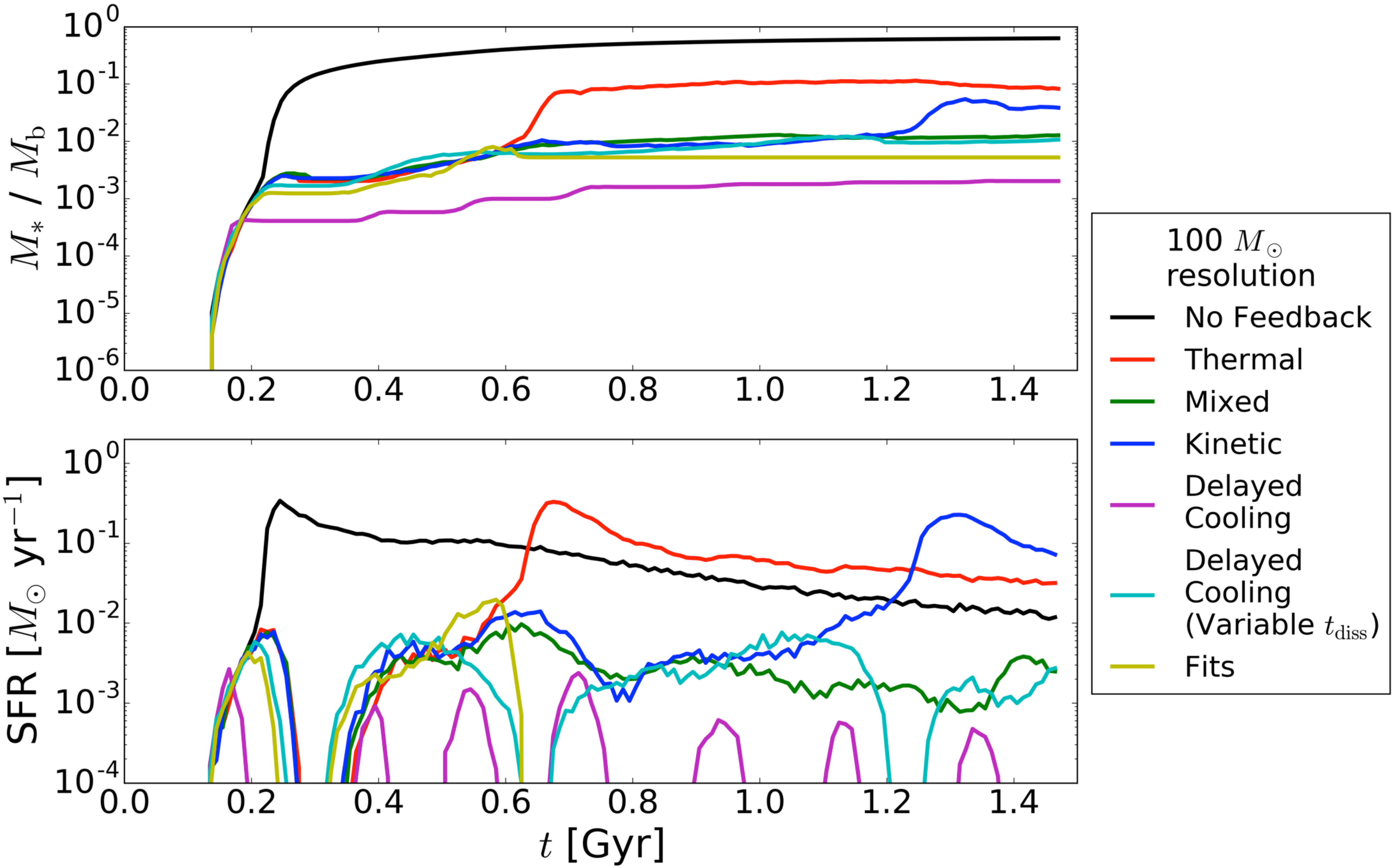


$M_* \,/\, M_{\rm b}$
SFR [$M_\odot$ yr^{-1}]
t [Gyr]
100 $M_\odot$ resolution
No Feedback
Thermal
Mixed
Kinetic
Delayed Cooling
Delayed Cooling (Variable $t_{\rm diss}$)

Fits to high resolution simulations

- Use fits from Martizzi+ 2015
- Fits to high resolution simulations of individual blastwaves
- Obtain thermal energy and momentum as a function of blastwave radius, background density and metallicity





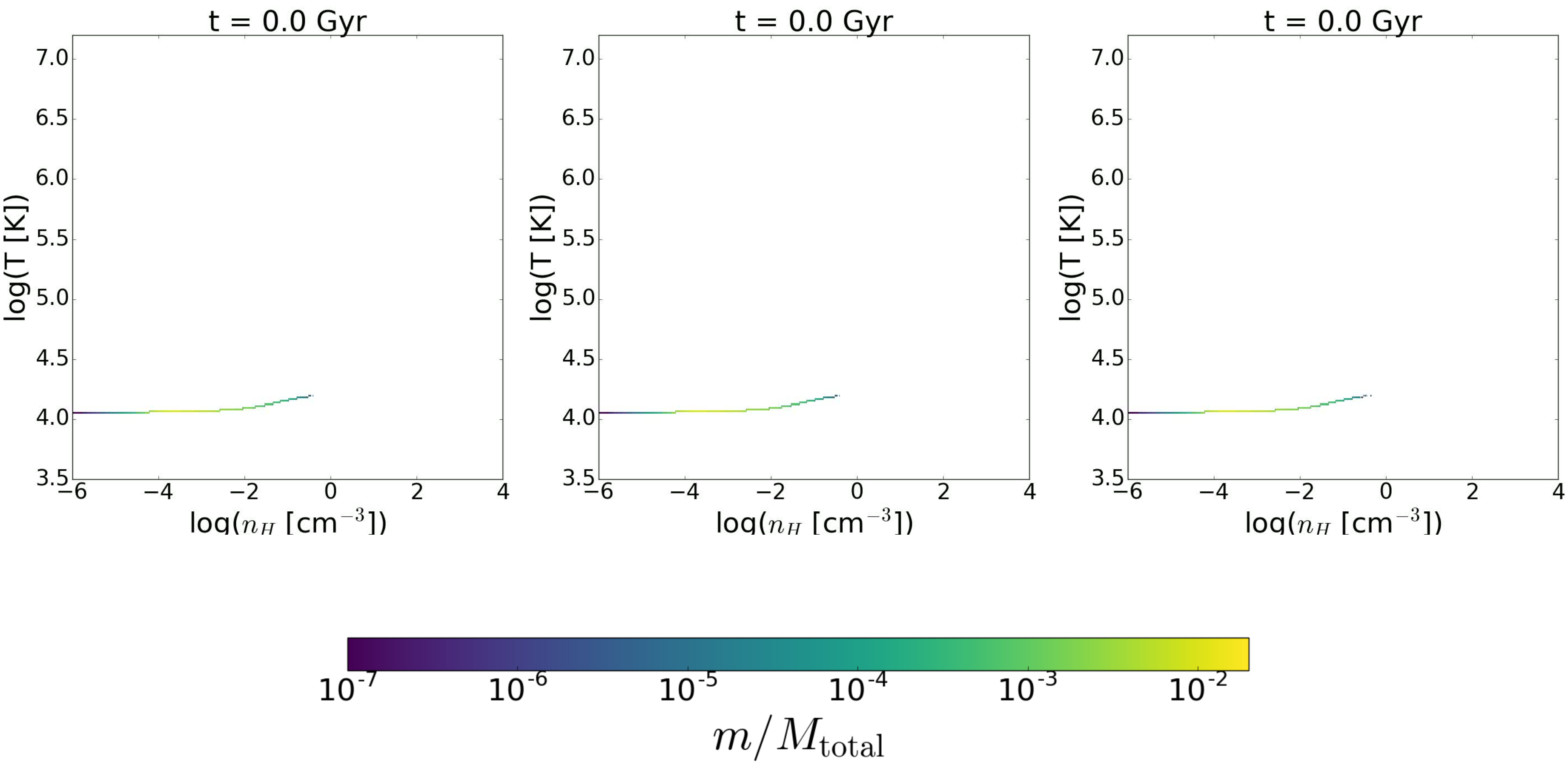
M_* / M_b
SFR [$M_\odot$ yr^{-1}]
t [Gyr]
100 $M_\odot$ resolution
No Feedback
Thermal
Mixed
Kinetic
Delayed Cooling
Delayed Cooling (Variable $t_{\rm diss}$)
Fits

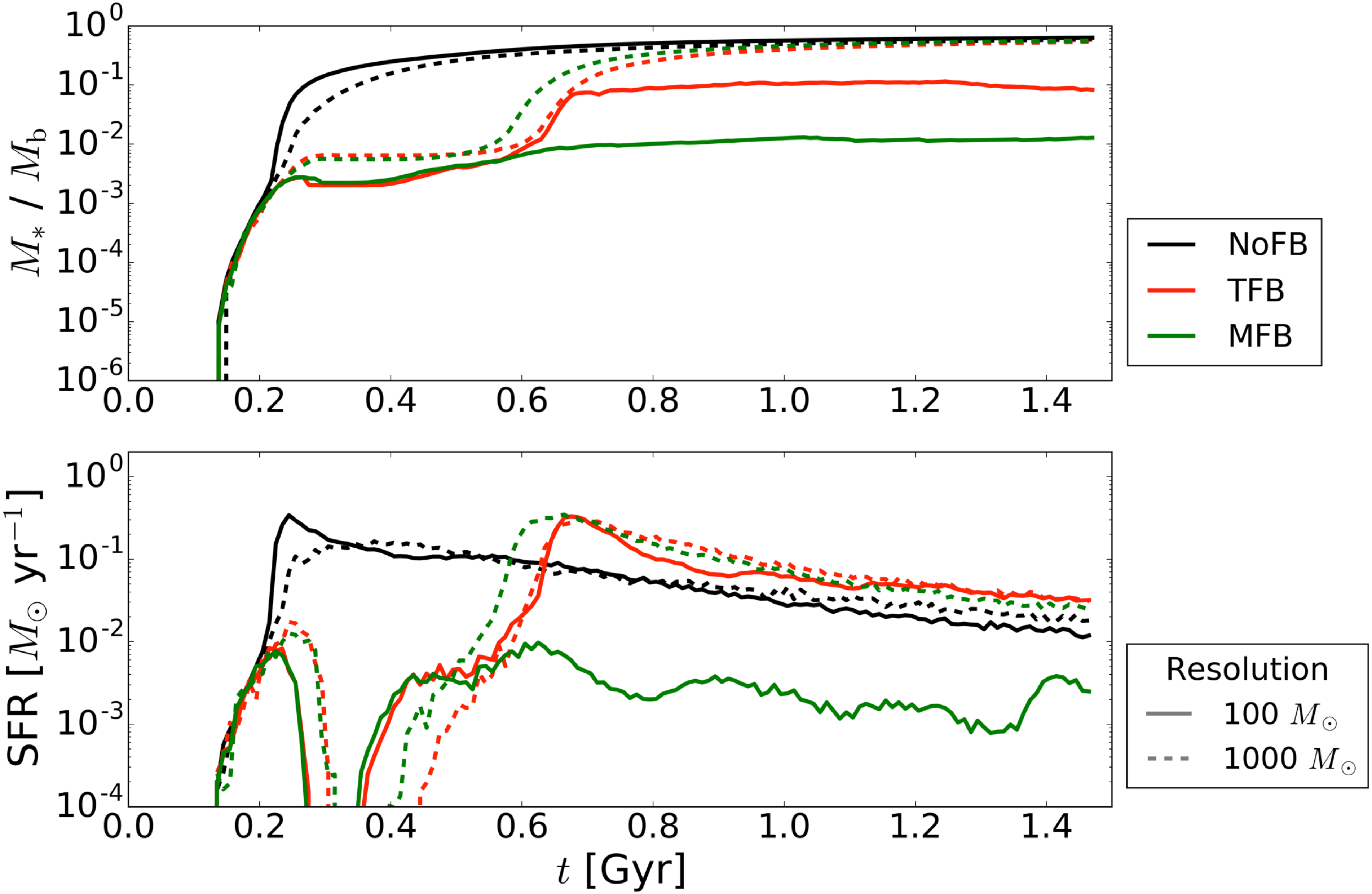
Phase Diagrams

Thermal

Delayed Cooling

Fits





$M_* \ / \ M_b$
10^0
10^{-1}
10^{-2}
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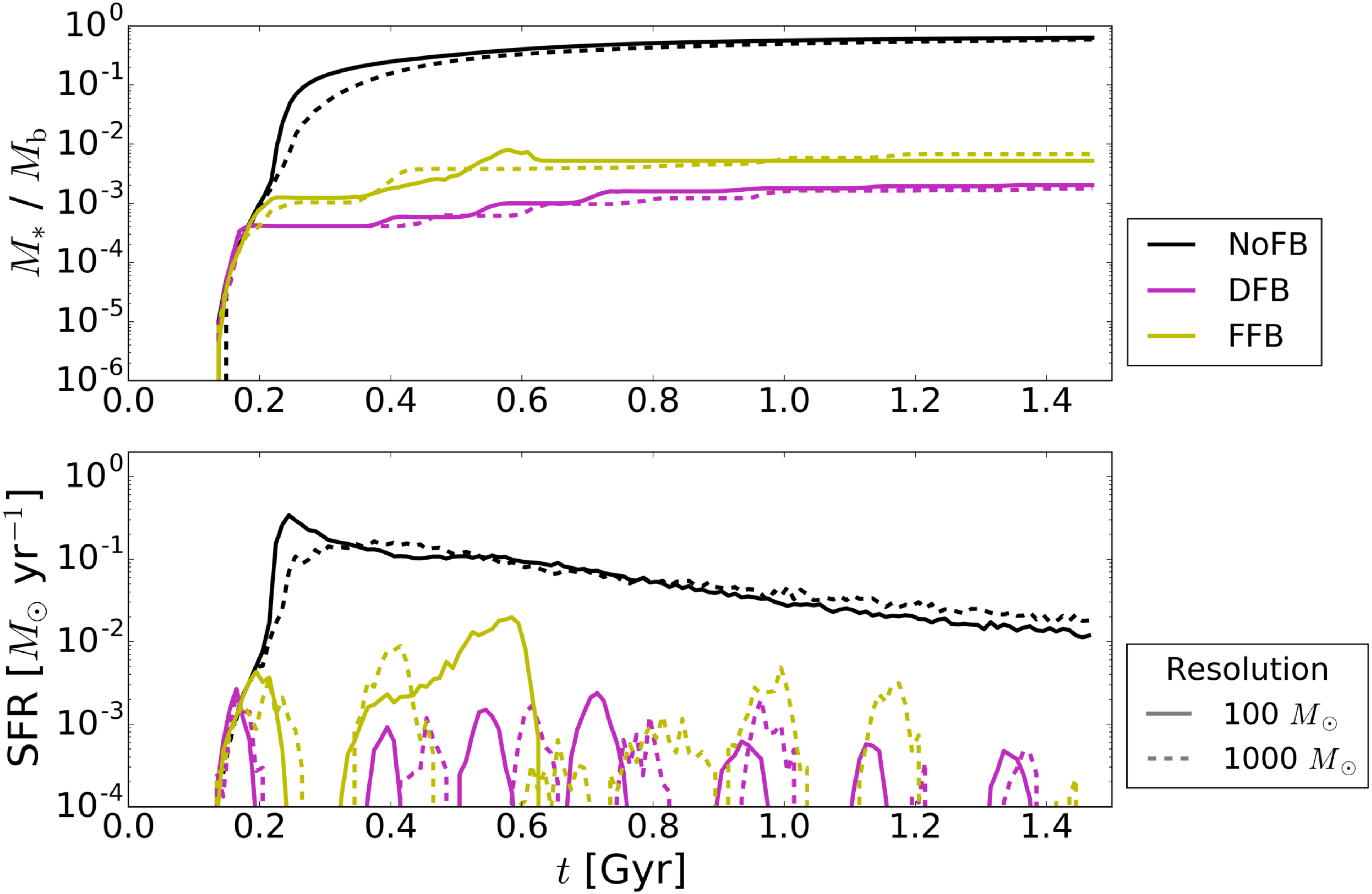
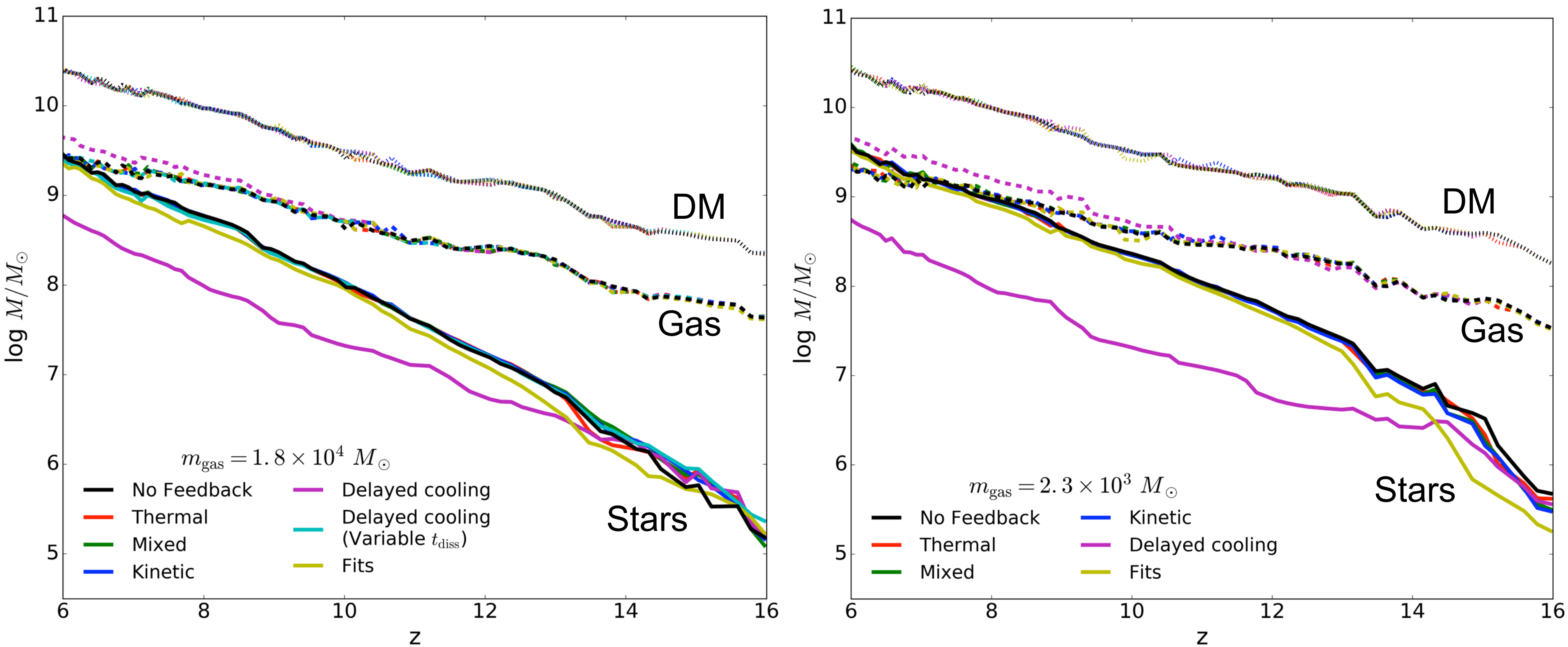
10^{-4}
10^{-5}
10^{-6}
0.0
0.2
0.4
0.6
0.8
1.0
1.2
1.4
NoFB
TFB
MFB
SFR $[M_\odot \ \mathrm{yr}^{-1}]$
10^0
10^{-1}
10^{-2}
10^{-3}
10^{-4}
Resolution
100 $M_\odot$
1000 $M_\odot$
t [Gyr]



$M_* \ / \ M_\mathrm{b}$
SFR [$M_\odot$ yr^{-1}]
t [Gyr]
NoFB
DFB
FFB
Resolution
100 $M_\odot$
1000 $M_\odot$

Cosmological zoom-ins



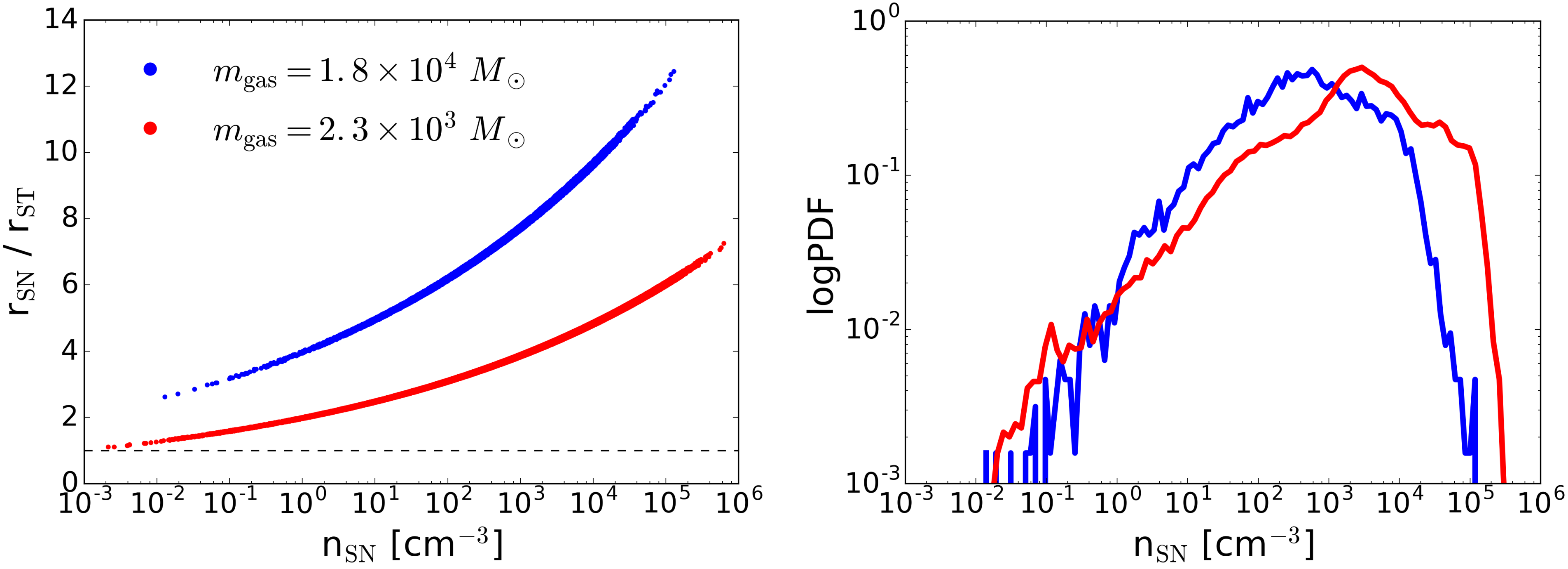
$m_{gas} = 1.8\times10^4\ M_\odot$
$m_{DM} = 6.7\times10^4\ M_\odot$

$m_{gas} = 2.3\times10^3\ M_\odot$
$m_{DM} = 8.3\times10^3\ M_\odot$

UV background (Haardt & Madau 2012)
Self-shielding (Rahmati+ 2013)

Supernovae sites

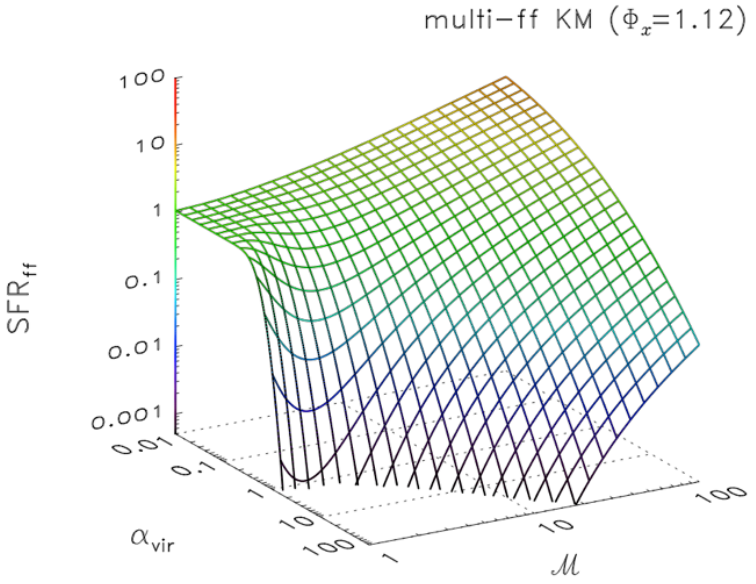
Thermal feedback



Non-SN processes

- 'Pretreatment' of ISM by other feedback processes will enhance SN effectiveness e.g.
  - Stellar winds
  - Radiation pressure
  - Photoionisation
- More realistic 'clustered' star formation criteria e.g. Krumholz & McKee (2005), Federrath & Klessen (2012)

$$\varepsilon \rightarrow \mathrm{SFR}_{\mathrm{ff}}\left(\alpha_{\mathrm{vir}}, \mathcal{M}, b, \beta\right)$$



Federrath & Klessen (2012)

Summary

- Supernova feedback is a key ISM process that must be included in cosmological simulations.
- Resolution issues make accurate modelling of SN feedback challenging.
- Results can be strongly dependent on choice of feedback implementation. Consequences of choice are non-trivial and must be understood in order to produce physically relevant simulations.
- SN feedback alone is not the whole picture; other processes and more physical star formation criteria will interact with SN feedback in a complex manner.